

## THE WHITE DWARF IN EM CYGNI: BEYOND THE VEIL\*

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Received 2009 February 19; accepted 2009 May 5; published 2009 June 19

### ABSTRACT

We present a spectral analysis of the *Far Ultraviolet Spectroscopic Explorer* (*FUSE*) spectra of the eclipsing double-line spectroscopic binary EM Cygni (EM Cyg), a Z Cam DN system. The *FUSE* spectrum, obtained in quiescence, consists of four individual exposures (orbits): two exposures, at orbital phases  $\phi \sim 0.65$  and  $\phi \sim 0.90$ , have a lower flux; and two exposures, at orbital phases  $\phi = 0.15$  and  $0.45$ , have a relatively higher flux. The change of flux level as a function of the orbital phase is consistent with the stream material (flowing over and below the disk from the hot spot region to smaller radii) partially masking the white dwarf. We carry out a spectral analysis of the *FUSE* data, obtained at phase 0.45 (when the flux is maximal), using synthetic spectra generated with the codes TLUSTY and SYNSPEC. Using a single white dwarf spectral component, we obtain a white dwarf temperature of  $40,000 \text{ K} \pm 1000 \text{ K}$ , rotating at  $100 \text{ km s}^{-1}$ . The white dwarf, or conceivably, the material overflowing the disk rim, shows suprasolar abundances of silicon, sulphur, and possibly nitrogen. Using a white dwarf+disk composite model, we obtain that the white dwarf temperature could be even as high as  $50,000 \text{ K}$ , contributing more than 90% of the FUV flux, and the disk contributing less than 10% must have a mass accretion rate reaching  $10^{-10} M_{\odot} \text{ yr}^{-1}$ . The single white dwarf model fits the absorption lines better than the white dwarf+disk model, but the white dwarf+disk model fits better the continuum in the shorter wavelengths. In both cases, however, we obtain that the white dwarf temperature is much higher than previously estimated. We emphasize the importance of modeling the spectra of EM Cyg around phase  $\phi < 0.5$ , when the white dwarf and disk are facing the observer, and we suggest that the discrepancy between the present analysis and previous spectral analysis might be due to the occulting effect of the stream veiling the white dwarf and disk.

*Key words:* accretion, accretion disks – novae, cataclysmic variables – stars: individual (EM Cyg) – ultraviolet: stars – white dwarfs

*Online-only material:* color figures

## 1. INTRODUCTION

### 1.1. Dwarf Novae

Dwarf novae (DNe) are a class of nonmagnetic cataclysmic variables (CVs) in which a white dwarf (WD) accretes matter from a main-sequence star (the *secondary*) filling its Roche lobe by means of an accretion disk. DN systems are characterized mainly by their brightness variations, due to a periodic change in the accretion rate. Ongoing accretion at a low rate (quiescence) is interrupted every few weeks to months by intense accretion (outburst) of days to weeks which increases the luminosity of the system (Hack & La Dous 1993; Warner 1995). This outburst is believed to be due to an instability in the disk (the disk instability model (DIM; Cannizzo 1998)), and it is known as a *dwarf nova* accretion event (or dwarf nova outburst).

Due to their long-term accretion, the WDs in DNe have temperatures  $\sim 15,000\text{--}50,000 \text{ K}$  (higher than field WDs), and therefore peak in the far-ultraviolet (FUV). The long periods of quiescence allow for long exposure observations ( $\sim$ hr) covering all the orbital phases (this is especially important for systems with a high inclination, as the WD can be eclipsed or masked at given orbital phases). Because of that, DNe, unlike

other CVs, are ideal candidates to study the accreting WD, as during quiescence the WD is exposed in the FUV for long periods of time. FUV spectroscopic observations are used to determine the WD parameters such as temperature, gravity, rotational velocity, and chemical abundances. The disk mass accretion rate of many systems can also be deduced accurately at given epochs of outburst or near outburst using spectral fitting techniques. In addition, DNe offer a fairly reliable estimate of their distances via the absolute magnitude at maximum versus orbital period relation (Warner 1995; Harrison et al. 2004). As a consequence, DNe are ideal candidates to observe, as one can more easily derive their system parameters, which in turn can be used to test the theories of CV evolution.

Dwarf nova systems are further divided into sub-types: U Gem systems, SU UMa systems, and Z Cam systems. The U Gem systems are the typical DNe, i.e., with normal DN outbursts; the SU UMas exhibit both normal DN outbursts, and superoutbursts, which are both longer in duration and higher in luminosity than normal DN outbursts; and the Z Cam systems have standstills where they remain in a semi-outburst state for a long time (Ritter & Kolb 2003). The binary orbital period in CV systems ranges from a fraction of an hour to about a day; however, there is a gap in the orbital period between 2 and 3 hr where almost no CV systems are found: the *period gap*. U Gem and Z Cam DN systems are found above the period gap, while the SU UMa DN systems are found below the period gap.

\* Based on observations made with the NASA-CNES-CSA *Far Ultraviolet Spectroscopic Explorer* (*FUSE*). *FUSE* is operated for NASA by the Johns Hopkins University under NASA contract NAS5-32985.

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# Report Documentation Page

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1. REPORT DATE <b>10 JUL 2009</b>		2. REPORT TYPE		3. DATES COVERED <b>00-00-2009 to 00-00-2009</b>	
4. TITLE AND SUBTITLE <b>The White Dwarf in EM Cygni: Beyond the Veil</b>				5a. CONTRACT NUMBER	
				5b. GRANT NUMBER	
				5c. PROGRAM ELEMENT NUMBER	
6. AUTHOR(S)				5d. PROJECT NUMBER	
				5e. TASK NUMBER	
				5f. WORK UNIT NUMBER	
7. PERFORMING ORGANIZATION NAME(S) AND ADDRESS(ES) <b>Department of Astronomy and Astrophysics, Villanova University, Villanova, PA, 19085</b>				8. PERFORMING ORGANIZATION REPORT NUMBER	
9. SPONSORING/MONITORING AGENCY NAME(S) AND ADDRESS(ES)				10. SPONSOR/MONITOR'S ACRONYM(S)	
				11. SPONSOR/MONITOR'S REPORT NUMBER(S)	
12. DISTRIBUTION/AVAILABILITY STATEMENT <b>Approved for public release; distribution unlimited</b>					
13. SUPPLEMENTARY NOTES					
14. ABSTRACT					
15. SUBJECT TERMS					
16. SECURITY CLASSIFICATION OF:			17. LIMITATION OF ABSTRACT	18. NUMBER OF PAGES	19a. NAME OF RESPONSIBLE PERSON
a. REPORT <b>unclassified</b>	b. ABSTRACT <b>unclassified</b>	c. THIS PAGE <b>unclassified</b>			

**Table 1**  
System Parameters for EM Cyg

Parameter	Value	References
$P_{orb}$	7 hr/0.29090912 days	Mumford & Krezminski (1969)
$i$	$67 \pm 2$ deg	North et al. (2000)
$M_{wd}$	$1 M_{\odot}$	Welsh et al. (2007); Robinson (1974); Stover et al. (1981)
$M_{2nd}$	$0.77 M_{\odot}$	Welsh et al. (2007); Robinson (1974); Stover et al. (1981)
$d$	350–500 pc	Bailey (1981); Jameson et al. (1981); Welsh et al. (2005)
$V_{min}$	14.4	Spogli et al. (2003, 2005)
$V_{max}$	12.5	Spogli et al. (2003, 2005)
DNO	14.6–16.5 s	Stiening et al. (1982)
2nd Spec. Type	K3V	North et al. (2000)
Arkadash Spec. Type	K2V–K5V	North et al. (2000)
Gal. Lat. $b$	$4^{\circ}27'$	SIMBAD
$E(B - V)$	$0.05 \pm 0.03$	Verbunt (1987); La Dous (1991); Bruch & Engel (1994)

In this paper, we present the FUV spectral analysis of the eclipsing Z Cam system EM Cygni observed during quiescence with *FUSE*.

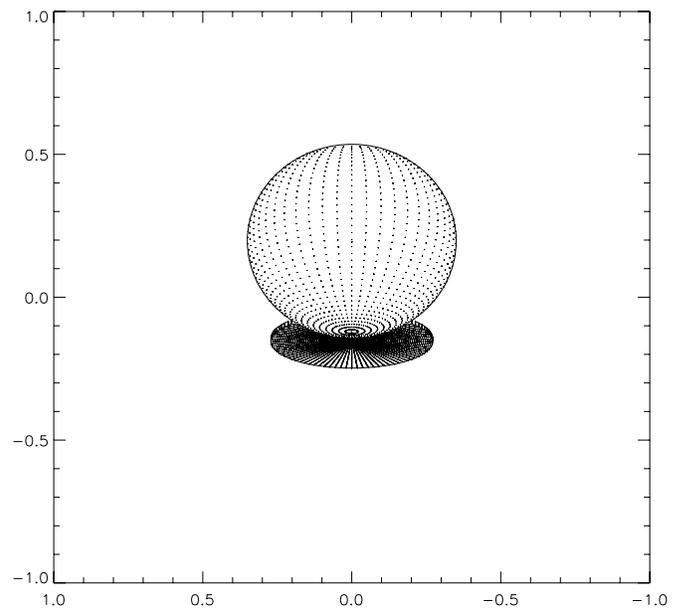
### 1.2. EM Cygni System Parameters

First suspected to be related to CVs by Burbidge & Burbidge (1953), EM Cyg is a bright, double-lined spectroscopic binary variable (Kraft 1964). Its eclipses were discovered by Mumford & Krezminski (1969) and the orbital period was found to be almost 7 hr (0.29090912 day). Its period was expected to decrease (by  $\sim 10^{-11}$ – $10^{-10}$  days/cycle; Pringle 1975; Mumford 1980), however it has been found to be constant (Beuermann & Pakull 1984; Csizmadia et al. 2008) with an upper limit  $\dot{P}/P < 2.3 \times 10^{-12}$  day $^{-1}$ . Its outburst cycle is of the order of 25 days and rapid coherent oscillations are detected during the late stages of outbursts with periods of the order of 14.6–16.5 s (see Stiening et al. 1982, and references therein).

EM Cyg is so far the only known Z Cam subtype eclipsing DN system. The eclipses are important as they allow, in theory at least, for a detailed study of the system. In addition, because EM Cyg is an eclipsing system as well as a double-lined spectroscopic binary, it has been extensively studied (e.g., Robinson 1974; Nevo & Sadeh 1978; Bailey 1981; Jameson et al. 1981; Patterson 1981; Stover et al. 1981; Szkody 1981; see also Stiening et al. 1982). However, the determination of the stellar masses of the system by Robinson (1974) and Stover et al. (1981) was effected by the presence of a third star, *Arkadash*,<sup>5</sup> along the line of sight (North et al. 2000). This third star is not related to the system. The most recent study (Welsh et al. 2007) indicates a mass ratio  $q = 0.77$  and a WD mass  $M_{wd} = 1 M_{\odot}$ . The distance found by Bailey (1981); Jameson et al. (1981),  $d \sim 320$ – $350$  pc, could also be affected by the contribution of *Arkadash* in the red light, and Welsh et al. (2005) pointed out that  $d$  could be as large as  $\sim 450$ – $500$  pc. However, using the maximum magnitude at outburst versus orbital period relation (Warner 1995; Harrison et al. 2004), one finds a distance of  $\sim 400$  pc (Winter & Sion 2003).

A recent broadband photometric study of the system (Spogli et al. 2003, 2005) shows that the magnitude of EM Cyg varies between  $V = 13.3$  at minimum and  $V = 12.5$  at maximum, it

<sup>5</sup> The third star was named *Arkadaş* (“friend” in Turkish) by Welsh et al. (2005), in memory of Janet (Hannula) Akyüz Mattei (1943–2004). Janet Mattei determined an outburst cycle of 25 days for EM Cyg, and provided AAVSO data and information on the state of EM Cyg many times (e.g., Stover et al. 1981; Stiening et al. 1982; North et al. 2000), as she did for many variable stars on countless occasions.

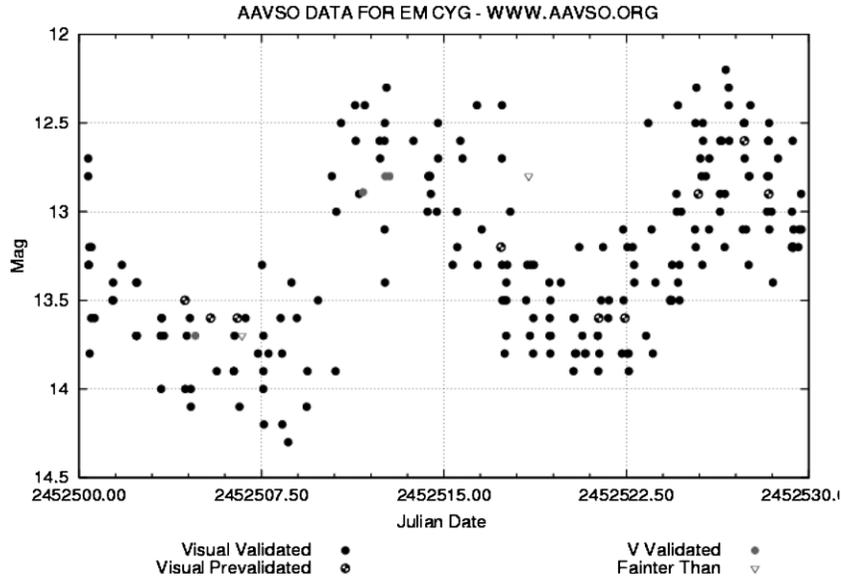


**Figure 1.** Binary configuration of EM Cyg is shown during eclipse of the disk at orbital phase  $\phi = 0$ . The inclination is  $i = 67^{\circ}$  and the mass ratio is 0.77. The disk is eclipsed but the white dwarf is not.

reaches  $V = 14.4$  during the eclipse and  $V = 12.9$  in standstill (note, however, that data from the AAVSO indicate that it can reach  $V \sim 12.0$ ). EM Cyg is located at a very low Galactic latitude ( $b = 4^{\circ}28'$ ), where the Galactic reddening reaches almost 0.5. However, the reddening toward EM Cyg is only  $0.05 \pm 0.03$  (Verbunt 1987; La Dous 1991; Bruch & Engel 1994). The system has an inclination  $i$  of only  $67 \pm 2$  deg (North et al. 2000), however, because of its large disk and secondary star, the disk is partially eclipsed, while the WD itself is not eclipsed. The geometrical configuration of the disk, WD and secondary of the system is shown in Figure 1 at phase  $\phi = 0/1.0$ . All the system parameters are recapitulated in Table 1.

### 1.3. Previous Spectral Analysis of EM Cyg

EM Cyg was observed in quiescence with *FUSE* by Welsh et al. (2005), who carried out a spectral fit after combining the *FUSE* spectrum with an *International Ultraviolet Explorer* (*IUE*) spectrum. Their spectral fit consists of a disk with  $\dot{M} = 2.5 \times 10^{-12} M_{\odot} \text{ yr}^{-1}$  and a  $0.88 M_{\odot}$  WD with  $T \approx 10$ – $20,000$  K, assuming a distance of 350 pc. The model did not fit the flux in the short wavelengths of *FUSE* which is indicative of a much higher temperature. However at  $T > 20,000$  K



**Figure 2.** Visual magnitude of EM Cyg recorded by the members of the AAVSO ([www.aavso.org](http://www.aavso.org)). EM Cyg was observed with STIS (around JD2452515) when the system was in outburst (center of the graph), and was observed with *FUSE* a week later (on JD2452523) when the system was in quiescence.

the model spectrum became too blue and the fit was noticeably worse in the *IUE* range.

Winter & Sion (2003) carried out a spectral fit to the *IUE* spectrum SWP 08088 of EM Cyg taken in quiescence (with a flux of  $\sim 3 \times 10^{-14}$  ergs cm $^{-2}$  Å $^{-1}$ ) and obtained that the best fit is a disk with  $\dot{M} = 5 \times 10^{-11} M_{\odot}$  yr $^{-1}$  and a WD of  $1 M_{\odot}$  with  $T < 24,000$  K (contributing only 8% of the flux) with a distance of 210 pc and  $i = 75^{\circ}$ . Later on, Urban & Sion (2006) analyzed the same *IUE* spectrum and obtained the same results except a distance of 350 pc, when dereddening the spectrum assuming  $E(B - V) = 0.05$ . Hamilton et al. (2007) analyzed the *IUE* spectrum SWP 07297 of EM Cyg in outburst assuming  $d = 350$  pc,  $i = 75^{\circ}$ , and obtained that the flux comes mostly from a disk with  $\dot{M} = 7 \times 10^{-9} M_{\odot}$  year $^{-1}$ .

In the spectral analysis of Winter & Sion (2003) and Urban & Sion (2006) only the *IUE* spectra of EM Cyg in quiescence was modeled. While in the work of Welsh et al. (2005) the fitting of the short wavelength range of *FUSE* (<1000 Å) was not carried out satisfactorily (as pointed out by Welsh et al. 2005 themselves). It is the purpose of this work to carry out a spectral fit to the *FUSE* spectrum of EM Cyg in quiescence including the short wavelength range, and to obtain the physical parameters of the exposed WD.

We give particular importance to the orbital phase at which the individual *FUSE* exposures were obtained. We show that at particular orbital phases the WD might be veiled by the stream from the first Lagrangian point overflowing the disk. We remark that previous estimates might have modeled a flux attenuated by as much as 75%, and consequently obtained a low WD temperature. We model here the *FUSE* spectrum of the WD when veiling is minimum and obtain a temperature  $T \sim 40,000$  K and possibly as high as 50,000 K.

In the next section, we present the four individual *FUSE* exposures and compare them with the *IUE* spectra taken in quiescence; in Section 3 we review the stream–disk interaction and in particular we check the possibility of the stream overflowing the disk; in Section 4, we describe the synthetic stellar spectral code as well as our fitting technique; the results are presented in Section 5, and a discussion with our conclusions are given in the last section.

## 2. OBSERVATIONS OF EM CYG

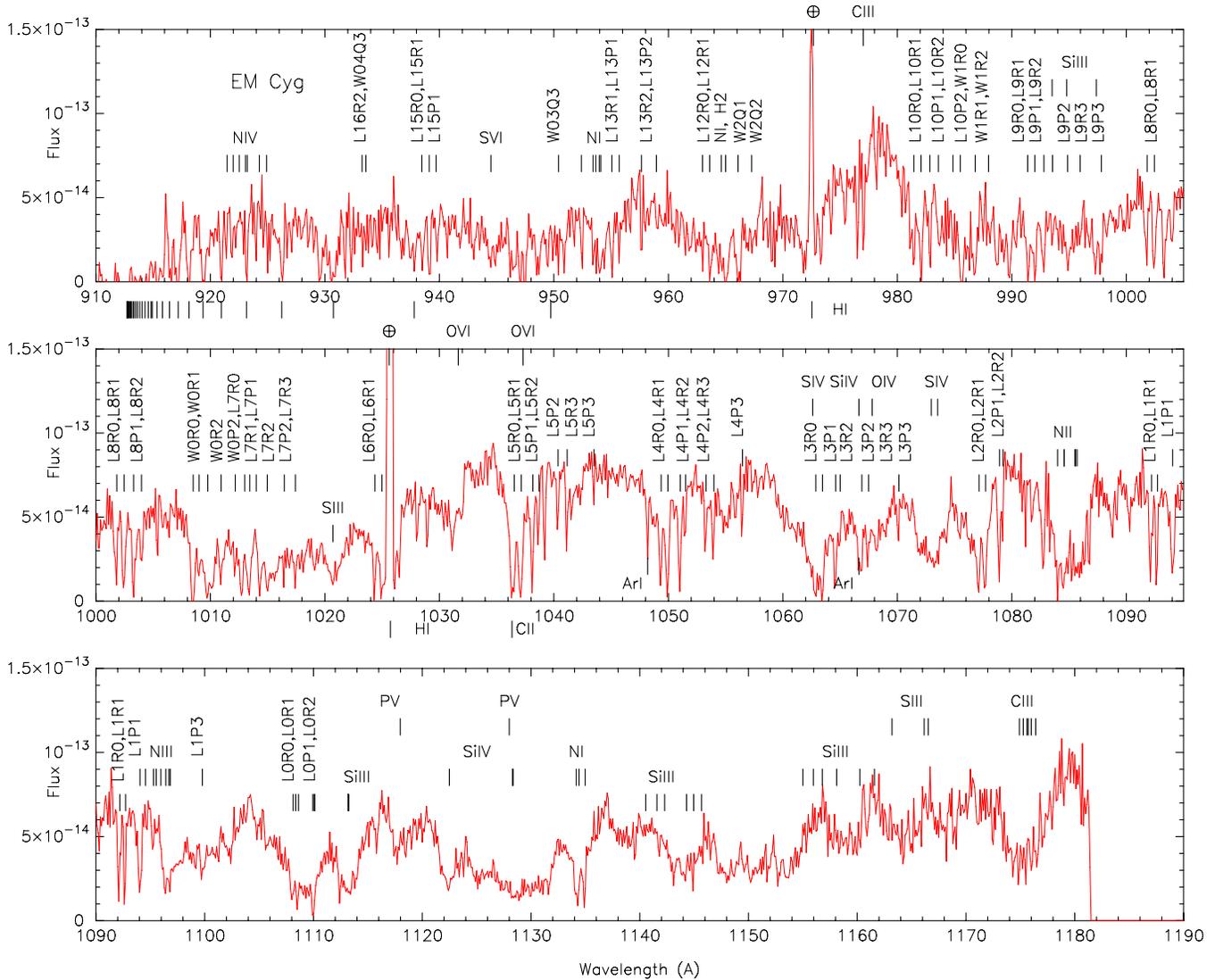
### 2.1. The *FUSE* Spectrum

The *FUSE* data of EM Cyg were obtained on 2002 September 5 (HJD2452522-3, when the system was in a low state, see Figure 2) through the  $30'' \times 30''$  LWRS Large Square Aperture in TIME TAG mode. The data were processed with the latest and final version of CalFUSE (v3.2.0; Dixon et al. 2007). We process the *FUSE* data as in Godon et al. (2008), and details of the procedure can be found there.

The *FUSE* spectrum of EM Cyg consists of four individual exposures (one per orbit). For the purpose of line identifications we present the averaged spectrum in Figure 3. We identify many interstellar medium (ISM) molecular hydrogen absorption lines as well as some of the metal lines usually seen in the *FUSE* spectra of accreting WDs. The metal lines we identify are as follows.

1. The S III (1021 Å), S IV (1062 & 1073 Å), S VI (945 Å) lines, and O VI doublet, all indicative of a higher temperature  $T \sim 35,000$  K and much higher.
2. The Si III (1114 Å), Si IV (1122 Å), and N II (1085 Å) lines, sometimes indicative of a lower temperature  $T \sim 25,000$  K, but not inconsistent with a higher temperature.
3. The C III (1175 Å) absorption line, present at all WD temperatures, as long as the carbon abundance is a fraction of its solar value.

The C II (1010 Å & 1066 Å) absorption lines, seen at lower temperatures ( $T \sim 25$ – $30,000$  K), are not detected here (though the C II (1010 Å) could be contaminated with ISM molecular absorption). There are some broad emission features from C III (977 Å), and the O VI doublet. The C III (1175 Å) presents a P-Cygni profile. The sharper absorption lines are from the ISM and possibly also from circumbinary material. It is likely that the nitrogen absorption lines (N I and N II) are contaminated with terrestrial absorption. All these lines are marked and identified in Figure 3. There are a few additional absorption features/lines that we cannot identify, mainly in the longer wavelengths (these lines are not labeled in the lower panel of Figure 3). The sharp emission lines are from air glow (geo- and helio-coronal



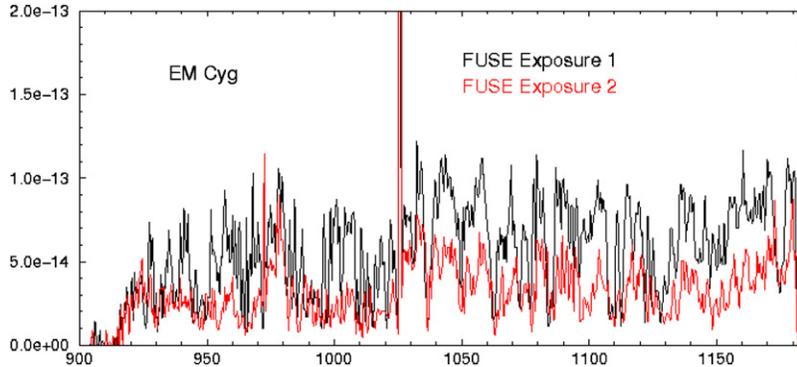
**Figure 3.** *FUSE* spectrum of EM Cyg. This spectrum consists of 4 exposures (corresponding to four *FUSE* orbits). The sharp emission lines are contamination. The ISM molecular features are marked vertically for clarity. The molecular lines are identified by their band (W = Werner; L = Lyman), upper vibrational level (1–16), and rotational transition (R, P, or Q with lower rotational state  $J = 1–3$ ). The ISM atomic lines are marked below the x-axis. Some ISM metal lines are also identified: N I, Ar I. The Nitrogen lines are possibly also contaminated with terrestrial atmospheric absorption (especially the N I 1135 Å). We have marked the position of the N IV 924 Å lines, though they are not detected. The C III and O VI lines are in broad emission, with a possible P-Cygni profile at the C III 1175 Å. (A color version of this figure is available in the online journal.)

in origin). The broad Ly $\beta$  absorption feature is not clearly seen, pointing to a high temperature, some broad oxygen doublet emission contamination, and/or possibly a rotational velocity broadening from a disk component. In the short wavelengths (<950 Å), the flux does not go to zero and the continuum there is about 1/3 of the flux at longer wavelengths (>1000 Å), indicative of a high temperature and consistent with the presence of higher-ionization metal lines.

As mentioned, this *FUSE* spectrum consists of four individual exposures taken during four consecutive orbits of the *FUSE* telescope. The four exposures correspond to the binary phases  $\phi \sim 0.45, 0.65, 0.90,$  and  $0.15$ , where  $\phi = 0$  corresponds to the eclipse (Welsh et al. 2005, 2007). The observations log for the individual *FUSE* exposures is given in Table 2. Using an analysis of the lines in the four individual *FUSE* exposures, Welsh et al. (2005, 2007) explicitly showed how the radial velocity changes as a function of the orbital phase. However, they did not consider the variation of the continuum flux level

from exposure to exposure, i.e., the variation of the FUV flux as a function of the orbital phase. In Figure 4, we show how the flux varies from exposure to exposure, and this change is also registered in the visual magnitude of the system as shown in data from the AAVSO in Figure 2. We find that at phases 0.15 and 0.45 (exposures 4 and 1), the *FUSE* continuum flux is larger (by up to a factor three) than at phases 0.65 and 0.90 (exposures 2 and 3).

In order to understand the variation of the FUV flux, in Figures 5 we show the binary configuration at the times the four *FUSE* exposures were obtained. From the orbital configuration we see that the flux in exposures 2 and 3 is attenuated when the hot spot region and the stream overflow region are in front of the observer. In the next subsection, we carry out an examination of the IUE spectra of EM Cyg, and we find that they too present some flux variation as a function of the orbital phase. We suspect that the FUV flux is possibly attenuated by absorption from the stream material (originating at L1) overflowing the disk edge



**Figure 4.** *FUSE* exposures 1 and 2 of EM Cyg. During the first exposure the flux is about 2 to 4 times larger than during exposure 2. Exposure 3 is similar to exposure 2; exposure 4 is similar to exposure 1. For clarity exposures 3 and 4 are not shown, and the spectra have been binned at 0.5 Å.

(A color version of this figure is available in the online journal.)

**Table 2**  
FUSE and IUE Observations Log

Obs.	Date (dd/mm/yy)	Time (UT)	Exp. Time (s)	Data Set	Start HJD	Phase <sup>a</sup> ( $\phi$ )	V (Range)
<i>FUSE</i>	05/09/02	11 <sup>h</sup> 34 <sup>m</sup> 11 <sup>s</sup>	1479	C0100101001	2452522.9849	0.42–0.48	13.1–13.9
<i>FUSE</i>	05/09/02	12 <sup>h</sup> 54 <sup>m</sup> 33 <sup>s</sup>	2651	C0100101002	2452523.0407	0.62–0.73	13.1–13.9
<i>FUSE</i>	05/09/02	14 <sup>h</sup> 34 <sup>m</sup> 31 <sup>s</sup>	2648	C0100101003	2452523.1102	0.85–0.96	13.1–13.9
<i>FUSE</i>	05/09/02	16 <sup>h</sup> 22 <sup>m</sup> 59 <sup>s</sup>	2135	C0100101004	2452523.1855	0.11–0.19	13.1–13.9
<i>IUE</i>	29/02/80	14 <sup>h</sup> 42 <sup>m</sup> 50 <sup>s</sup>	2700	LWR 07056	2444299.1102	0.86–0.98	13.6–13.7
<i>IUE</i>	29/02/80	15 <sup>h</sup> 32 <sup>m</sup> 22 <sup>s</sup>	4500	SWP 08088	2444299.1445	0.98–0.16	13.6–13.7
<i>IUE</i>	05/08/82	02 <sup>h</sup> 40 <sup>m</sup> 47 <sup>s</sup>	1500	LWR 13862	2445186.6152	0.65–0.71	13.2–13.5
<i>IUE</i>	05/08/82	03 <sup>h</sup> 11 <sup>m</sup> 30 <sup>s</sup>	1500	SWP 17591	2445186.6366	0.72–0.78	13.2–13.5
<i>IUE</i>	05/08/82	03 <sup>h</sup> 44 <sup>m</sup> 26 <sup>s</sup>	2400	LWR 13863	2445186.6595	0.81–0.91	13.2–13.5
<i>IUE</i>	05/08/82	04 <sup>h</sup> 34 <sup>m</sup> 08 <sup>s</sup>	2399	SWP 17592	2445186.6940	0.92–0.02	13.2–13.5

**Note.**

<sup>a</sup> The ephemeris, time of minima/eclipse HJD2400000 + 43780.7508, 45257.4091, 50692.4613, & 53709.1873 were taken from Mumford (1980); Beuermann & Pakull (1984); Csizmadia et al. (2008) with the period 0.29090912 days = 25134.55 s.

and reaching a height above (and below) the disk of the order of  $z/r \sim 0.4$ . We further check this suggestion in Section 3.

### 2.2. The IUE Spectra

Since the *FUSE* flux varies as a function of the orbital phase it is interesting to compare it with other spectra obtained in quiescence; this will also help us interpret results from previous spectral analysis. The existing archival *Hubble Space Telescope*/Space Telescope Imaging Spectrograph (*HST*/STIS) spectra of EM Cyg were taken during outburst, just a week before the acquisition of the quiescent *FUSE* spectrum (see Figure 2), and the three existing HUT spectra were also taken in outburst. Because of that we disregard the STIS and HUT spectra. However, some of the existing *IUE* spectra were taken in quiescence, and later were modeled by Winter & Sion (2003); Welsh et al. (2005); Urban & Sion (2006). We present and discuss these spectra here.

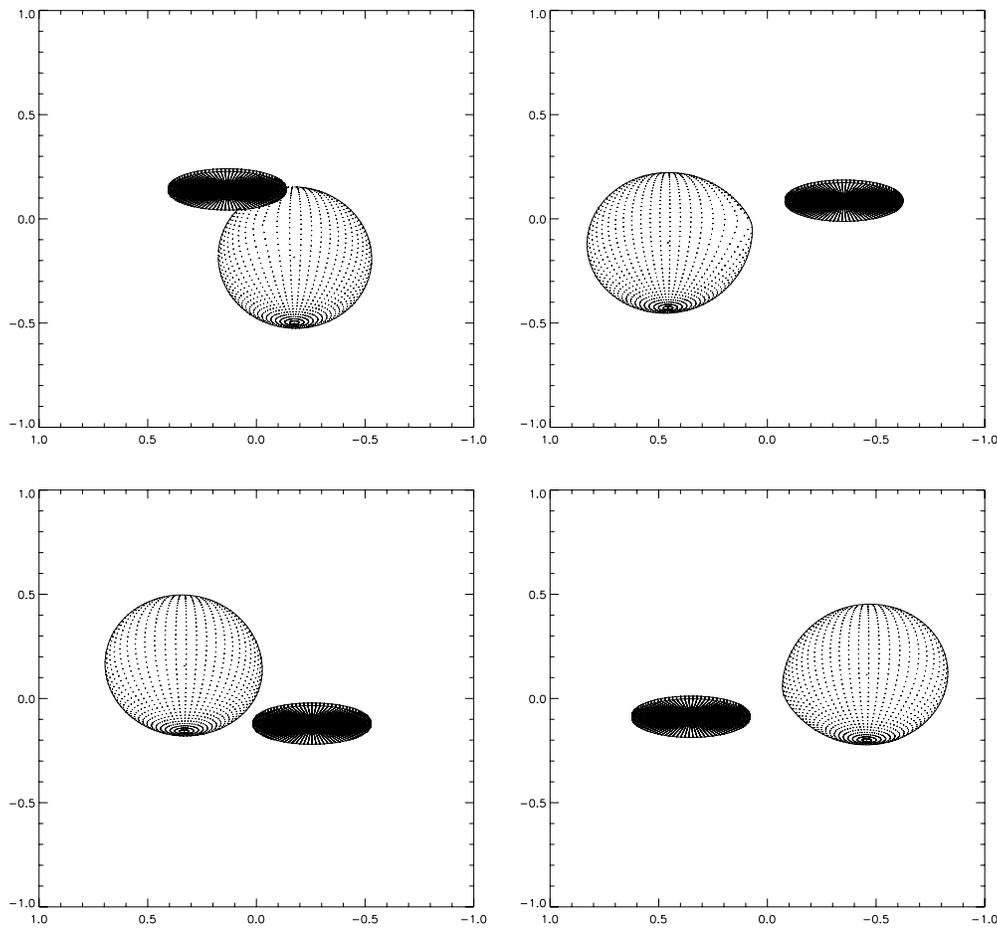
In Figure 6, we compare three *IUE* SWP spectra taken in quiescence together with the *FUSE* exposures in the spectral range  $\sim 900$ – $1,400$  Å. Since *FUSE* exposures 2 and 3 have about the same flux, we combine them together, and the same for exposures 1 and 4. We see that SWP 08088 (taken in 1980) has a much lower flux, similar to exposures 2 and 3. SWP 17591 (taken in 1982) has a flux level comparable with exposures 1 and 4, while SWP 17592 (taken 90 minutes after SWP 17591) has a higher flux level.

The typical quiescent *V* of the system is 13.5 but it has been observed as faint as 14.5 in quiescence (possibly at eclipse

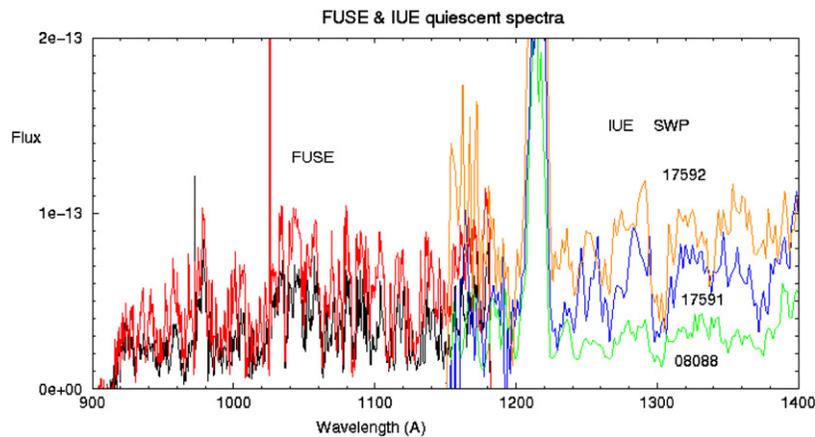
Spogli et al. 2003, 2005). From validated AAVSO data, it appears that, at the time of the *FUSE* observations, the system had a magnitude of  $V \sim 13.1$ – $13.9$  (Figure 2), namely EM Cyg was not in its lowest possible quiescent state at the time of the *FUSE* observations. This is clearly seen in Figure 2, where the previous quiescent state on JD2452505 is about 0.5 mag fainter than the quiescent state during which the *FUSE* observations (JD2452522) were carried out. The *IUE* SWP 17591 and SWP 17592 spectra were obtained during a brighter state, when EM Cyg was on the rise to outburst, its magnitude around 13.2–13.5 (but still lower than a standstill with  $V \sim 12.5$ – $13.5$ ). The *IUE* SWP 08088 spectrum was at  $V \sim 13.6$ – $13.7$  (Szkody 1981).

The three *IUE* SWP spectra are complemented with three longer wavelength (LWR) spectra. All the quiescent (LWR and SWP) *IUE* spectra are listed in Table 2 together with their corresponding orbital phase. The increase in flux between SWP 17591 and SWP 17592 over a time scale of 90 minutes is most probably due to orbital variation (as in *FUSE*), and the matching LWR spectra 13862 and 13863 do not show any variation. As to the SWP 08088 and (matching) LWR 07056 spectra, their lower flux is due to a deeper quiescence and to orbital variation too.

Because of these changes in flux level (due to both orbital variation and intrinsic changes in the source) combining the *FUSE* and *IUE* spectra together is tricky and this might explain the difficulty Welsh et al. (2005) were having to match both the short wavelengths of *FUSE* and the longer wavelengths of *IUE* in the same fit. As a consequence, we do not combine the



**Figure 5.** EM Cyg Binary Configuration during the four *FUSE* orbits/exposures. Exp.1 (upper left) is at phase  $\phi \sim 0.45$ : the disk is in front of the observer and absorption is minimal. Exp.2 (upper right) is at phase  $\phi \sim 0.65$ : the disk is still slightly in front moving away from the observer, the stream overflow is in front, absorption is maximal. Exp.3 (lower left) is at phase  $\phi \sim 0.90$ : the secondary filling its Roche lobe is in front moving to the right, the stream and hot spot are directly in front, absorption is maximal. Exp.4 (lower right) is at phase  $\phi \sim 0.15$ : the secondary Roche lobe is slightly in front and to the right of the observer, the disk is in the back moving toward the observer. The stream is in the back and behind, absorption is minimal.



**Figure 6.** *FUSE* exposures are shown together with the SWP *IUE* spectra. The *FUSE* exposures 1 and 4 have been combined together and are shown in red; exposures 2 and 3 have also been combined together and are shown in black. *FUSE* exposures 2 and 3 have a lower flux due to veiling of the WD by stream material overflowing the disk. It is likely that the lower flux in the *IUE* spectrum SWP 08088 is also due to the veiling of the WD. Additional change in the *IUE* spectra might be due to the cooling of the WD or ongoing accretion.

(A color version of this figure is available in the online journal.)

*FUSE* spectrum of EM Cyg with any *IUE* spectrum. The shorter wavelength range as well as the higher signal-to-noise ratio (S/N) of *FUSE* over *IUE* makes the *FUSE* spectrum of EM Cyg ideal for determining the temperature of the accreting WD.

From the *FUSE* & *IUE* spectra of EM Cyg, we found that the FUV flux between  $\sim 900 \text{ \AA}$  and  $2000 \text{ \AA}$  is attenuated around phase  $\phi \sim 0.6\text{--}0.9$ . Since the *FUSE* spectrum presents many deep absorption lines, it is tempting to suggest the possibility of an absorbing component moving in front of the observer around

phase  $\sim 0.75$ , as the binary rotates. We further investigate this possibility in the next section.

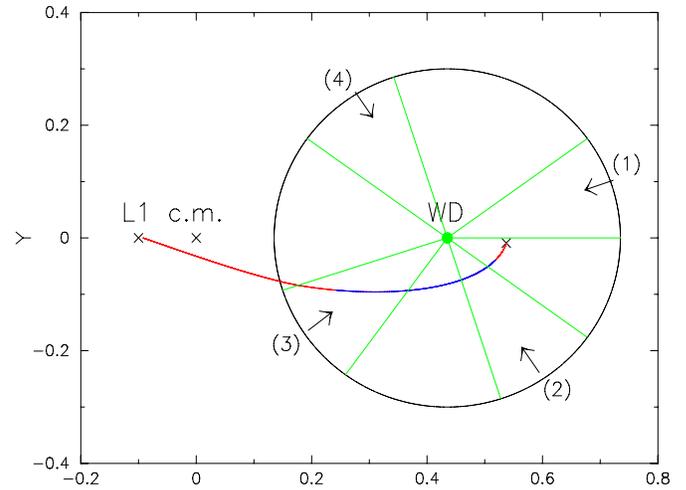
### 3. THE STREAM-DISK INTERACTION

The only components of the system that could be in the field of view of the WD/inner disk around  $\phi \sim 0.6-0.9$  are the region downstream the hot spot and the stream possibly overflowing disk at the hot spot. In order to check this possibility, in this section, we consider previous studies of the stream-disk interaction, and also present a simple ballistic trajectory simulation using the binary parameters of EM Cyg. We show that for the binary parameters of EM Cyg, stream material could be found high enough above the disk of the plane, around phase  $\phi \sim 0.6-0.9$  to actually veil the WD and inner disk.

Pioneering analytical studies by Lubow & Shu (1976); Lubow (1989) have shown how part of the stream material possibly flows over and under the disk and falls back onto the disk at smaller radii at an angle  $\phi_{\text{impact}} \approx 140^\circ-150^\circ$  ( $\phi_{\text{impact}}$  is measured counter clockwise from the line joining the two stars), corresponding to an orbital phase  $\phi \sim 0.6$ . Observationally, Doppler tomography (Marsh 1985) has been used for a quarter of a century to confirm the stream-disk interaction and the impact at smaller radii in some system (e.g., Z Cha). In Lubow (1989) some of the stream matter was flowing over the disk at a scale height  $\sim 0.01a$ , where  $a$  is the binary separation. For a disk of diameter  $R_{\text{disk}} = 0.3a$ , this corresponds to a scale height of  $z/r \approx 0.03$  at the edge of the disk. Namely, in these analytical studies, it was obtained that stream material overflows the disk rim (at the hot spot) and continue toward smaller radii while basically *grazing* the disk's surfaces (at a few vertical scale lengths—a few  $\sim H$ ). Therefore, based on these studies alone, in order for the material to veil the WD, the inclination of the system has to be  $\cos i \approx 0.03 \rightarrow i \approx 88^\circ$  or larger: the system has to be nearly edge-on.

However, more recent full three-dimensional hydrodynamical simulations, e.g., (Blondin 1998; Kunze et al. 2001; Bisikalo et al. 2003), have followed in detail the interaction of the material streaming from the first Lagrangian point (L1) with the disk. The main result from these simulations is that the stream does not simply grazed the disk's surfaces, but matter is actually deflected vertically from the hot spot and flows in a more or less diffuse stream to inner parts of the disk, hitting the disk surface close to the *circularization radius*  $r_c$  (see Lubow & Shu 1975 for details on  $r_c$ ) at orbital phase 0.5. In these simulations, the stream material deflected vertically at the hot spot region can reach an altitude  $z \sim r$ . This stream deflection is believed to cause X-ray absorption in CVs (and LMXBs) around orbital phase 0.7, if the inclination is at least  $65^\circ$  (Kunze et al. 2001). These simulations also explain streams of material perpendicular to the orbital plane as observed in Algol-type binaries (e.g., U CrB; Agafonov et al. 2009, and references therein). These results imply that stream material can be deflected vertically, and reaches a height  $z/r = \cos i = \cos 65^\circ = 0.42$  and lands onto the disk at  $\phi_{\text{impact}} = 180^\circ$  (corresponding to an orbital phase 0.5; Kunze et al. 2001).

We suggest the same might be happening in EM Cyg, and the decreased flux around  $\phi \sim 0.7$  might be caused by the deflected stream flow. In order to verify the plausibility of this scenario, we simulate the stream flow from the L1 point using the ballistic trajectory approximation, and we deflect the ballistic particle at the hot spot by an angle  $\theta$  (measure from the plane of the disk). Since the stream flow consists of a infinite number (in the continuity limit) of particles (and since the stream has

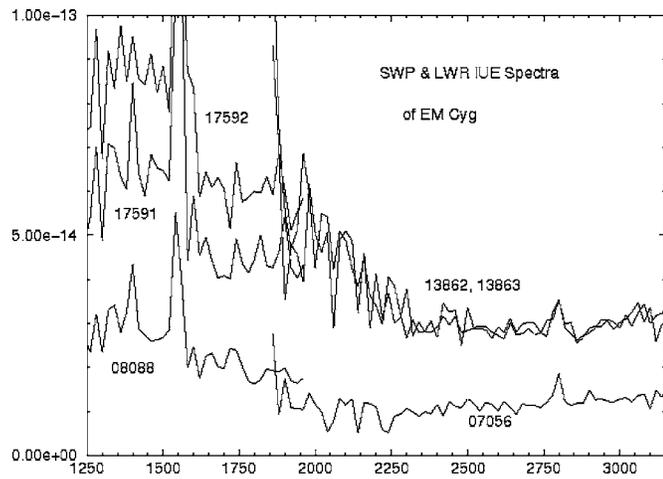


**Figure 7.** EM Cyg Binary Configuration during the four *FUSE* orbits/exposures. The areas delimited in green represent the viewing angles during the four *FUSE* exposures (1, 2, 3, & 4). The disk radius here is  $0.3a$ , where  $a$  is the binary separation. The stream trajectory from the L1 point is shown in red. The stream overflows the disk rim and from there it is launched at an angle  $\sim 45^\circ$  above the disk; it falls back onto the disk around  $\phi \sim 0.5$ . The region for which the stream can reach  $z/r \sim 0.4$  (corresponding to veiling at  $i \sim 67^\circ$ ) is shown in blue (from P. Godon 2009, in preparation).

(A color version of this figure is available in the online journal.)

a finite thickness extending vertically), the ballistic trajectory is solved for all possible value of  $\theta$  (from  $-90^\circ$  to  $+90^\circ$ ). A maximum altitude is reached for a critical angle of deflection  $\theta = \theta_{\text{Cri}}$ . This maximum altitude depends on the values of the system parameters: veiling will not occur for every mass ratio, disk radius, and inclination (the full details of the simulations as well as results for the entire space parameters of CVs are given elsewhere; P. Godon 2009, in preparation). Figure 7 depicts such a simulation for EM Cyg, where we show with arrows the four *FUSE* exposures. The stream is shown in red and the blue area represents that portion of the trajectory for which the material reaches  $z/r \geq 0.4$ , corresponding to veiling at  $i = 67^\circ$ . Here, the stream is deflected at the hot spot with an angle  $\theta_{\text{Cri}} \approx 45^\circ$  above (and below) the plane of the disk. This is the angle for which the altitude of the ballistic particle reaches a maximum. For a smaller, or *larger*, deflection angle  $\theta$ , the blue region shrinks. For a smaller disk radius the blue region shrinks too as the hot spot moves counterclockwise toward  $\phi \sim 0.75$ . This points to the importance of the large disk in EM Cyg for deflecting material off the disk plane around  $\phi \sim 0.6-0.9$ . This picture is consistent with the full three-dimensional hydrodynamical simulations of Kunze et al. (2001). For different system parameters, (inclination, mass ratio, and disk radius) the deflected material does not especially veil the WD, as the maximum vertical elevation of the ballistic particle does not always reach  $z/r = \cos i$ .

From this specific ballistic trajectory simulation with the system parameters of EM Cyg, we conclude that the lower flux in the *FUSE* exposures 2 and 3 can be accounted for with stream material being deflected vertically at the hot spot forming a veil in front of the WD between orbital phases  $\sim 0.65$  and  $\sim 0.90$  (for an inclination of  $67^\circ$ ). The increase in flux between the *IUE* spectra SWP 17591 ( $\phi \sim 0.75$ ) and SWP 17592 ( $\phi \sim 0.97$ ) over a time scale of 90 minutes can also be explained by veiling of the WD/inner disk by the stream overflow around phase 0.75 (keeping in mind that the WD/inner disk are not eclipsed around  $\phi \sim 1.0$ ).



**Figure 8.** The IUE LWR and SWP spectra of EM Cyg. The wavelength is in Å, and the flux is in  $\text{erg s}^{-1}\text{cm}^{-2}\text{Å}^{-1}$ . The plateau between 2250 and 3000 Å is possibly due to the hot spot with  $T \sim 10,000$  K, as pointed out by Szkody (1981).

Furthermore, the stream material overflowing the disk edge lands onto the disk at a smaller radius corresponding to orbital phase of up to  $\phi \sim 0.5$ , and produces a bright spot there (Marsh 1985; Lubow 1989). Together with the hot spot, this contributes to additional flux in the LWR spectra. The emission from the hot spot and point of impact is possibly seen in the LWR spectra (13862, 13863) shown in Figure 8 as a plateau in the longer wavelengths ( $\lambda \sim 2250\text{--}3000$  Å), where one rather expects the flux to decrease there. This flux in the longer wavelengths was already associated with emission from the hot spot by Szkody (1981). Though we have no way to confirm that the plateau in the longer wavelengths of IUE is due in part to the impact point of the stream overflow, it is, however, consistent with this possibility. The plateau seen in the LWR spectra, is also contributing significantly to the longer wavelengths of the SWP spectra. This is especially true for the SWP 08088 spectrum. These are additional reasons not to model the IUE spectra, nor to combine them with the existing *FUSE* spectrum, as the hot spot and point of impact contribute flux to the longer wavelengths of the IUE SWP spectra.

## 4. SYNTHETIC SPECTRAL MODELING

### 4.1. Modeling the Stellar Spectrum

We create model spectra for high-gravity stellar atmospheres using the codes TLUSTY and SYNSPEC<sup>6</sup> (Hubeny 1988; Hubeny & Lanz 1995). We generate photospheric models for a  $1.0 M_{\odot}$  WD with effective temperatures ranging from 12,000 K to 75,000 K in increments of about 10% (e.g., 1,000 K for  $T \approx 15,000$  K and 5,000 K for  $T \approx 70,000$  K). We vary the stellar rotational velocity  $V_{\text{rot}} \sin(i)$  from  $50 \text{ km s}^{-1}$  to  $500 \text{ km s}^{-1}$  in steps of  $10\text{--}50 \text{ km s}^{-1}$ .

We also use disk models from the grid of models generated by Wade & Hubeny (1998). We select their models with a WD mass  $M = 1 M_{\odot}$  and inclinations  $i = 60^{\circ}$  &  $75^{\circ}$ . These models have solar abundances, but because of the velocity broadening the absorption lines are completely smoothed out, such that the disk models ( $i \sim 60^{\circ}\text{--}75^{\circ}$ ) are not sensitive to the chemical abundances. These disk models are either used as such or in combination with WD models we generate.

After having generated grids of models, we use FIT (Press et al. 1992), a  $\chi^2$  minimization routine, to compute the reduced  $\chi^2_v$  ( $\chi^2$  per number of degrees of freedom) and scale factor values for each model fit. Before carrying out each synthetic spectral fit of the spectrum, we masked portions of the spectrum with strong emission lines and air glow. While we use a  $\chi^2$  minimization technique, we do not blindly select the least  $\chi^2$  models, but we also examine the models that best fit some of the features such as absorption lines.

The WD rotation ( $V_{\text{rot}} \sin(i)$ ) rate is determined by fitting the WD model to the spectrum while paying careful attention to the line profiles. We did not carry out separate fits to individual lines but rather tried to fit the lines and continuum in the same fit. We first fit solar abundance models and only then we change the abundances of the species.

### 4.2. Modeling the ISM Hydrogen Absorption Lines

We identify ISM lines in Figure 3 to avoid confusing them with the WD lines. The ISM lines are deep and broad and we decided to model them, especially since some of the WD lines (such as S iv  $\lambda\lambda 1062.6$  and  $1073$ ) are located at almost the same wavelengths.

We model the ISM hydrogen absorption lines to assess the atomic and molecular column densities. This enables us to improve the WD spectral fit. The ISM models (transition tables) are generated using a program developed by P. E. Barrett. This program uses a custom spectral fitting package to estimate the temperature and density of the interstellar absorption lines of atomic and molecular hydrogen. The ISM model assumes that the temperature, bulk velocity, and turbulent velocity of the medium are the same for all atomic and molecular species, whereas the densities of atomic and molecular hydrogen, and the ratios of deuterium to hydrogen and metals (including helium) to hydrogen can be adjusted independently. The model uses atomic data of Morton (2000, 2003) and molecular data of Abgrall et al. (2000). The optical depth calculations of molecular hydrogen have been checked against those of McCandliss (2003).

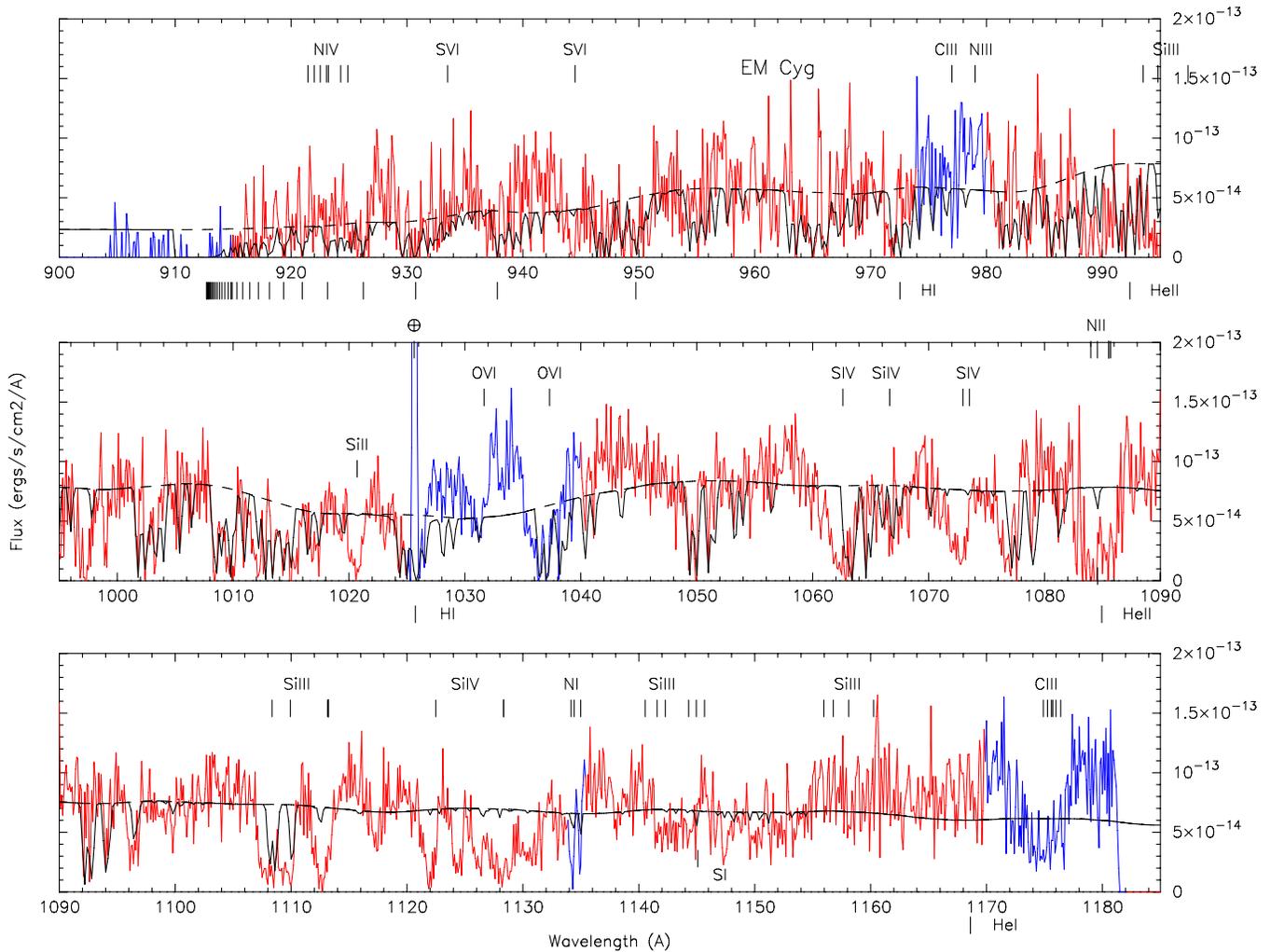
The ratios of metals to hydrogen and deuterium to hydrogen are fixed at 0 and  $2 \times 10^{-5}$ , respectively, because of the low S/N data. The wings of the atomic lines are used to estimate the density of atomic hydrogen and the depth of the unsaturated molecular lines for molecular hydrogen. The temperature and turbulent velocity of the medium are primarily determined from the lines of molecular hydrogen when the ISM temperatures are  $< 250$  K.

The ISM absorption features are best modeled and displayed when the theoretical ISM model (transmission values) is combined with a synthetic spectrum for the object (namely a WD synthetic spectrum or a WD+disk composite spectrum).

## 5. RESULTS

Though the distance estimate to EM Cyg is most like around 400 pc (Bailey 1981; Jameson et al. 1981; Winter & Sion 2003; Welsh et al. 2005), we do not fix the distance in our model fitting. We adopt a WD mass of  $1 M_{\odot}$  (Welsh et al. 2007) in all our models, with a corresponding radius of  $\sim 6000$  km (giving  $\log(g) \sim 8.6$ ). All the model fits are carried out on the first exposure (with the maximum flux). The reddening toward EM Cyg is  $0.05 \pm 0.03$ . As a consequence, we model EM Cyg assuming both  $E(B - V) = 0.0$  and  $E(B - V) = 0.05$ . We first carry out the fitting without dereddening the spectrum, and only later do we deredden the spectrum (assuming  $E(B - V) = 0.05$ ) to assess the effect of reddening on the results.

<sup>6</sup> <http://nova.astro.umd.edu>; TLUSTY version 200, SYNSPEC version 48.



**Figure 9.** Best-fit single disk model (in solid black) to the *FUSE*/exposure 1 spectrum of EM Cyg. The regions that have been masked for the fitting are shown in blue. The model consists of an accretion disk around a  $1 M_{\odot}$  WD with  $\dot{M} = 1 \times 10^{-9} M_{\odot} \text{ yr}^{-1}$ , an inclination  $i = 75^{\circ}$ . The distance obtained is  $d = 441$  pc and  $\chi^2_{\nu} = 0.6916$ . Because of the rotational broadening the model cannot fit the absorption lines from the source, and the mass accretion rate is too large and inconsistent with the quiescent state of the system. An ISM model has been added to the disk model to fit the ISM lines. The dotted line show the disk model without the inclusion of the ISM model.

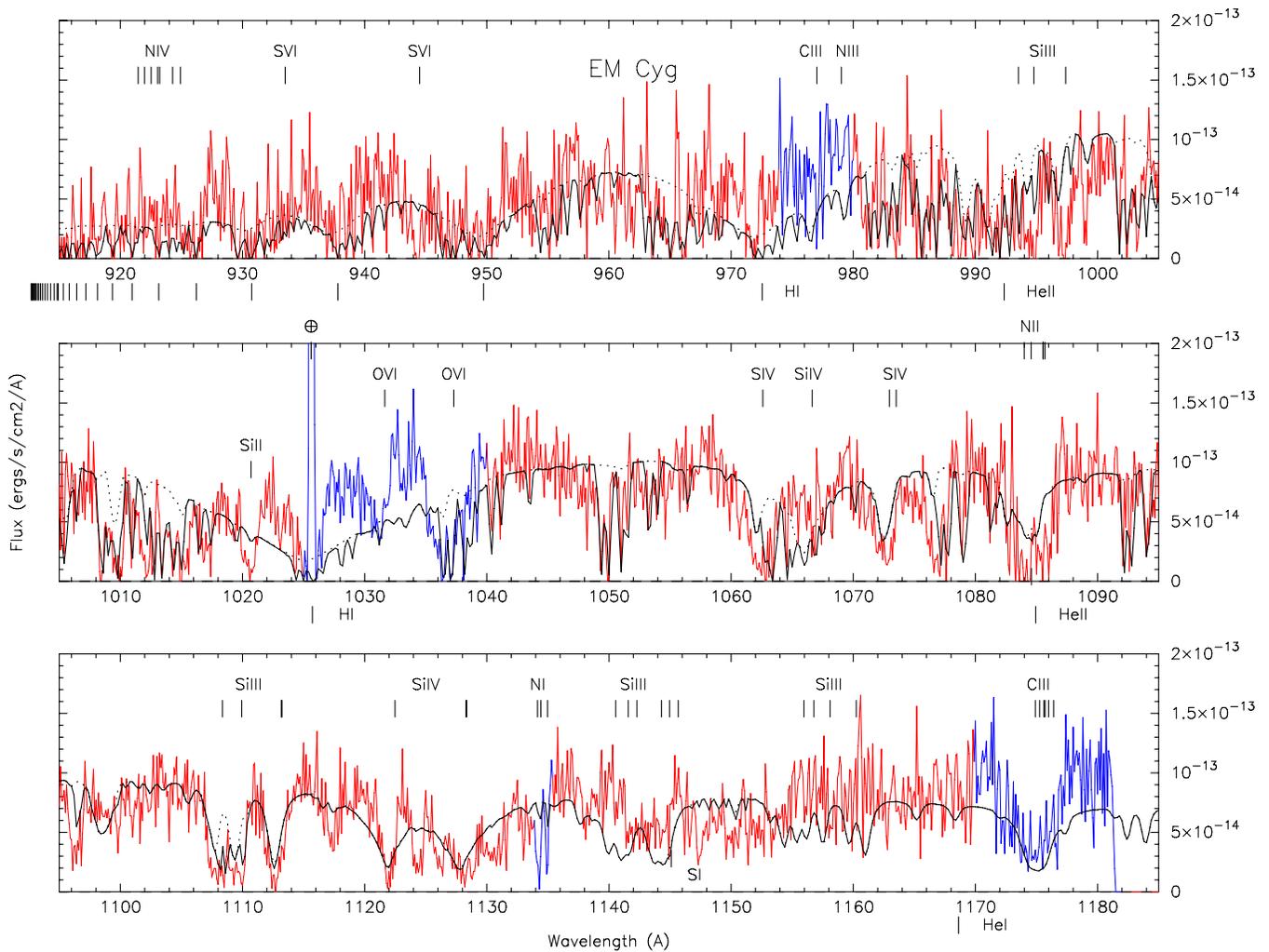
(A color version of this figure is available in the online journal.)

We first run single disk models. We find that in order to fit the short wavelengths of *FUSE*, the best single disk model has a mass accretion rate  $\dot{M} \sim 10^{-9} M_{\odot} \text{ yr}^{-1}$ , i.e., far larger than expected for quiescence, and  $i \sim 60^{\circ}$ – $75^{\circ}$ . The  $i = 75^{\circ}$  model is shown in Figure 9. It gives a distance of 441 pc and  $\chi^2_{\nu} = 0.6916$  (all our  $\chi^2$  values are smaller than one due to the low S/N of the *FUSE* exposures). On top of the model, we add an ISM model in order to optimize the fit. The ISM has the following parameters  $N_H = 2 \times 10^{19} \text{ atoms cm}^{-2}$ ,  $N_{H_2} = 1 \times 10^{17} \text{ molecules cm}^{-2}$ ,  $T = 160$  K,  $v = 25 \text{ km s}^{-1}$ ,  $D/H = 1 \times 10^{-5}$ , and  $Z = 0 \times$  solar. We did not try to optimize the ISM model itself, nor to obtain accurate parameter values for the ISM. The modeling of the ISM is carried only in the context of improving the disk/WD model fit and to help identify the lines and spectral features. It is clear from Figure 9 that this disk model only vaguely follows the continuum of the spectrum and does not fit the absorption features (except for the ISM lines). The two best single disk models are listed in Table 3.

Next, we run single WD models. We find that in order to fit the continuum the WD must have a temperature of 40,000 K, with a corresponding distance of 280 pc. The  $\chi^2$  obtained is

$\chi^2_{\nu} \approx 0.58$ . The two first WD models in Table 3 show how the distance and  $\chi^2$  change as a function of the WD temperature. In this first set we chose a guessed rotational velocity of  $200 \text{ km s}^{-1}$ . In the next model (Table 3), we improve the  $\chi^2$  value by varying the rotational velocity. The optimal value is somewhere between 90 and  $100 \text{ km s}^{-1}$ , and we decide to fix the rotational velocity to  $100 \text{ km s}^{-1}$ .

In the next four models (Table 3), we change the abundances of silicon, sulphur, and nitrogen while keeping all the other abundances to their solar value (we found that a solar carbon abundance is the best fit, though the main carbon line at  $1175 \text{ \AA}$  has also some broad emission and makes it difficult to model). The best fit is obtained for enhanced abundances as follows: sulphur: 10 times solar, nitrogen: 10 times solar, and silicon: 20–30 times solar. We obtain a distance less than anticipated:  $\sim 260$  pc. We present this best-fit model in Figure 10. In order to optimize the fit we shift the synthetic WD spectrum to the blue by  $0.5 \text{ \AA}$  to account for the radial velocity shift. This improves slightly the  $\chi^2$  values (the values listed in Table 3 reflect this improvement). We also add an ISM model to the WD model, however, the ISM model is not shifted.



**Figure 10.** Best-fit WD model (in solid black) to the *FUSE*/exposure 1 spectrum of EM Cyg. The observed spectrum (in red) has not been dereddened. The model consists of a  $1.0 M_{\odot}$  WD, with  $T = 40,000$  K, a projected rotational velocity  $V_{\text{rot}} \sin i = 100 \text{ km s}^{-1}$ , and over-solar abundances of Si (30), S (10), and N (10). The entire WD synthetic spectrum has been shifted to the blue by  $\sim 0.5 \text{ \AA}$ , which corresponds to a radial velocity of about  $\sim 150 \text{ km s}^{-1}$  (at  $1,000 \text{ \AA}$ ) in rough agreement with the radial velocity  $+125 \text{ km s}^{-1}$  Welsh et al. (2005). An ISM model (no shifted) has been added to improve the overall fit. The dotted line represents the WD model without the inclusion of the ISM model. The regions of the observed spectrum that have been masked for the fitting are shown in blue. The distance obtained from the best fit is  $d = 261 \text{ pc}$  and  $\chi^2_{\nu} = 0.4768$

(A color version of this figure is available in the online journal.)

In the next model, in Table 3, we deredden the spectrum assuming  $E(B - V) = 0.05$ , and obtain the same results, except for the distance. Here the distance is barely  $200 \text{ pc}$ . We note that assuming a WD mass of  $\sim 0.8 M_{\odot}$  can increase the distance to  $350 \text{ pc}$  without changing the remaining parameters of the fit. On the other side, adding an accretion disk does increase the distance too, and this is what we do next.

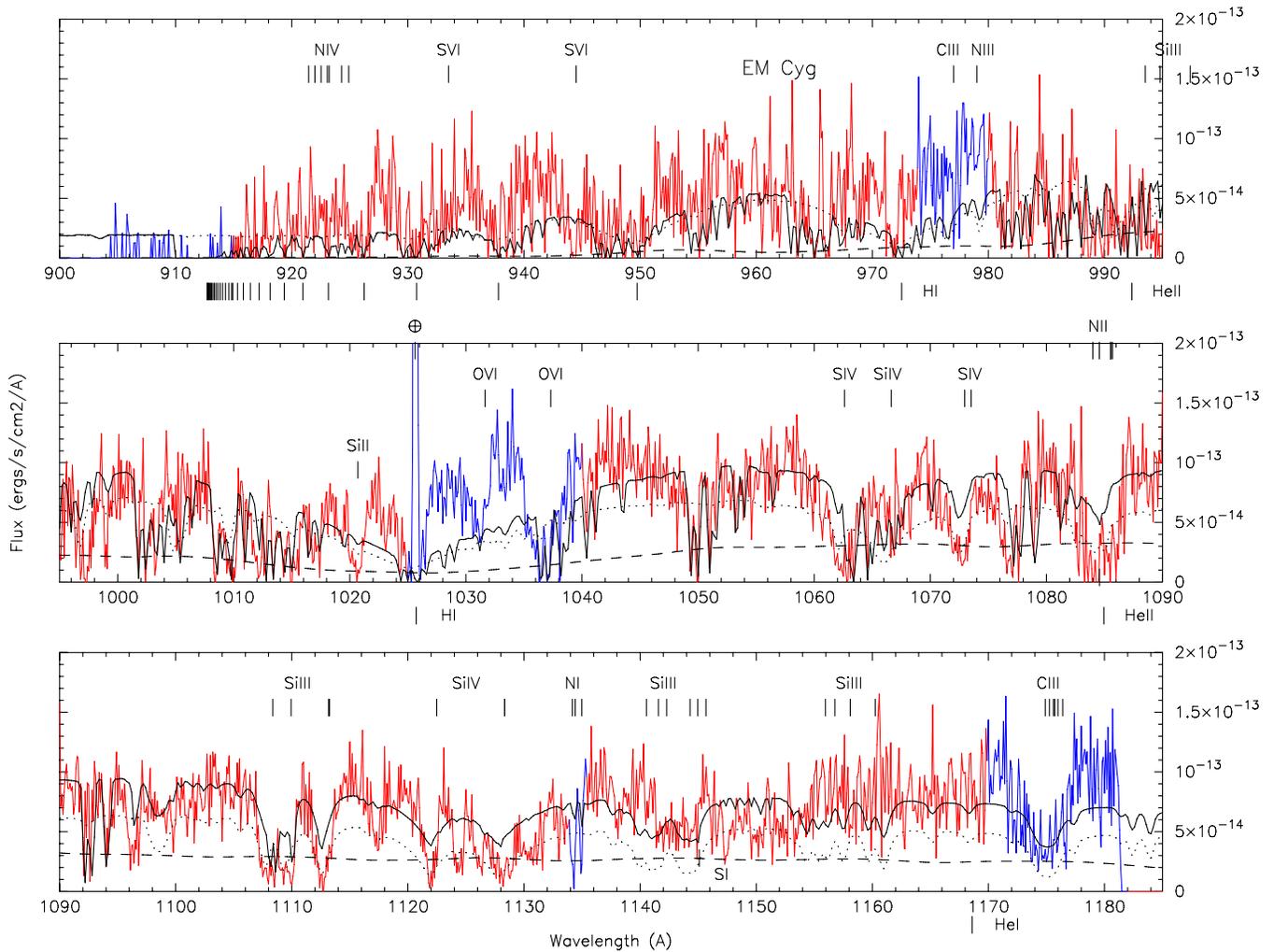
For the WD+disk composite model, we first take our best single WD model and add a disk model to it, to try and improve the fit and increase the distance. We find that the mass accretion rate has to be at least  $\dot{M} = 10^{-10} M_{\odot} \text{ yr}^{-1}$  for the disk to contribute a noticeable flux. As the disk is added, the model deteriorates and give a larger  $\chi^2_{\nu}$ . We present such a model in Figure 11. The fit in the lower wavelength is not as good as the best-fit single WD model presented in Figure 10.

In order to try and improve the fit in the shorter wavelength we increase the WD temperature, as the mass accretion rate is already too large for quiescence and therefore cannot be increased. We find that the fit in the shorter wavelengths can be improved by elevating the WD to a temperature  $T \approx 50,000 \text{ K}$ , however the overall  $\chi^2_{\nu}$  is further increased by doing so. This

model is shown in Figure 12. While the model does well at fitting the overall continuum of the spectrum, it does not fit the *FUSE* absorption lines, as the velocity broadening in the disk makes the overall depth of the line shallower in the WD+disk model. It is possible that the absorption lines form above the disk and WD; we discuss this possibility and its implications in the next and last section.

## 6. DISCUSSION AND CONCLUSIONS

The over abundances and the unidentified lines in the *FUSE* spectrum of EM Cyg might be associated with some of the veiling material rather than the WD itself, we have no additional way to verify this. However, simulations of the stream-disk interaction (Kunze et al. 2001) show that the stream overflow hits the disk at smaller radii where it can in part bounce back. If this is the case, stream material can be found at any orbital phase, but with greater density and altitude above the disk plane around  $\phi \sim 0.6-0.9$ . The rich variety of lines observed in the present *FUSE* spectrum of EM Cyg is reminiscent of the *HST*/STIS spectrum of the DN TT Crt observed in a low state with



**Figure 11.** Best WD model (from Figure 10) plus a disk model are fitted (in solid black) to the *FUSE*/exposure 1 spectrum of EM Cyg. The observed spectrum (in red) has not been dereddened. The regions that have been masked for the fitting are shown in blue. The disk model (dashed line) consists of an accretion disk around a  $1 M_{\odot}$  WD with  $\dot{M} = 1 \times 10^{-10} M_{\odot} \text{ yr}^{-1}$ , an inclination  $i = 60^{\circ}$ . The WD model (dotted line) consists of a  $1.0 M_{\odot}$  WD, with  $T = 40,000$  K, a projected rotational velocity  $V_{\text{rot}} \sin i = 100 \text{ km s}^{-1}$ , and over-solar abundances of Si (30), S (10), and N (10). The entire WD+disk synthetic spectrum has been shifted to the blue by  $\sim 0.5 \text{ \AA}$ . The distance obtained is  $d = 318 \text{ pc}$  and  $\chi^2_{\nu} = 0.5291$ . An ISM model has been added to the WD+disk model to fit the ISM lines. The WD contributes 77% of the FUV flux and the disk contributes the remaining 23%.

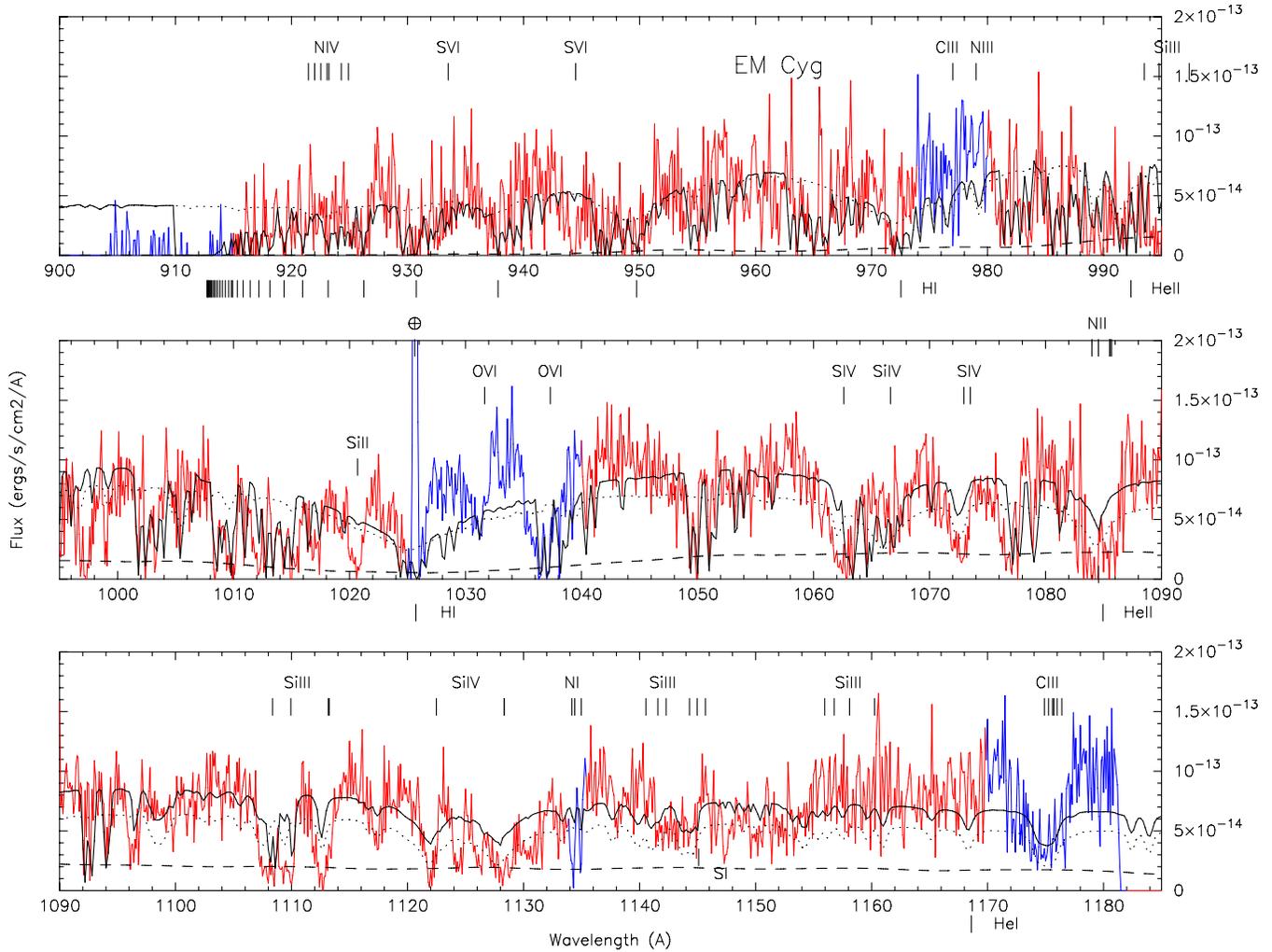
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an inclination of  $i \approx 60^{\circ}$  (Sion et al. 2008), and of the *FUSE* spectrum of the NL BB Dor possibly with an inclination of  $80^{\circ}$  (Godon et al. 2008). These spectra suggest line formation in different temperature regions, whether due to a wind above the disk (as could be the case in BB Dor) or overflowing stream material (possibly for TT Crt). So the strong and deep absorption lines in the *FUSE* exposure 1 of EM Cyg are likely due to circumstellar material, i.e., stream material overflowing the disk rim and bouncing off the disk at smaller radii, while the lower flux in exposures 2 and 3 are due to the direct veiling of the WD by the stream overflow around phase 0.7.

If the veiling material is responsible for the absorption lines, then the abundance fitting is not valid anymore, as the lines might be out of thermodynamic equilibrium or represent highly variable optical depths. As a consequence, the best fit to the continuum is actually the WD+disk model shown in Figure 12, rather than the single WD model. While the single WD model has a lower  $\chi^2$  value, the WD+disk model is in better agreement with the distance of the system and its higher-than-quiescent state. Because of this, we conclude that the *FUSE* spectrum of

EM Cyg is best interpreted as the spectrum of a hot WD with a small contribution from a low mass accretion disk.

The picture that emerges from this spectral analysis, is that of a 50,000 K WD, contributing 93% of the total FUV flux, and a disk, with  $\dot{M} = 10^{-10} M_{\odot} \text{ yr}^{-1}$  contributing the remaining 7%. Such a composite WD+disk model gives a distance  $d = 382 \text{ pc}$  in best agreement with its expected value. The WD+disk model is also in agreement with the moderately higher-than-quiescent state during which the *FUSE* spectrum was taken, even though the  $\chi^2$  value is slightly larger as the model does not fit the absorption lines very well. The spectrum presents some deep absorption lines as well as unidentified lines in the longer wavelength region ( $>1100 \text{ \AA}$ ), apparently forming in front and above the WD/disk, from stream material overflowing the disk edge. As a consequence, the exposure taken at phase  $\phi \sim 0.45$  has a higher flux than exposures obtained around  $\phi \sim 0.65\text{--}0.90$ . The large disk ( $0.3a$ ) in EM Cyg is consistent with this picture, as the hot spot is consequently located closer to the secondary (around  $\phi \sim 0.9$ ) and L1 stream material can more easily be found off the disk plane. The large disk



**Figure 12.** Same as in Figure 11, but here the WD temperature has been increased to 50,000 K. The WD contributes to 93% of the FUV flux and the disk contributes the remaining 7%. The distance obtained is  $d = 382$  pc and  $\chi^2_v = 0.5573$ . This model fits the shorter wavelength continuum flux better than the previous models, but it does not fit the absorption features as well as the single WD model.

(A color version of this figure is available in the online journal.)

**Table 3**  
WD Synthetic Spectra

$T_{wd}$ ( $10^3$ K)	$V_{rot} \sin i$ ( $\text{km s}^{-1}$ )	[Si] (Solar)	[S] (Solar)	[N] (Solar)	$\dot{M}$ ( $M_{\odot} \text{ yr}^{-1}$ )	$i$ (deg)	$d$ (pc)	$\chi^2_v$	WD/Disk (%)	$E_{B-V}$	Figure
...	...	1.0	1.0	1.0	$10^{-9}$	75	441	0.6916	0/100	0.00	9
...	...	1.0	1.0	1.0	$10^{-9}$	60	754	0.6867	0/100	0.00	
35	200	1.0	1.0	1.0	...	...	223	0.6063	100/0	0.00	
40	200	1.0	1.0	1.0	...	...	280	0.5823	100/0	0.00	
40	100	1.0	1.0	1.0	...	...	280	0.5709	100/0	0.00	
40	100	30.0	10.0	1.0	...	...	264	0.4780	100/0	0.00	
40	100	10.0	10.0	10.0	...	...	267	0.4847	100/0	0.00	
40	100	20.0	10.0	10.0	...	...	264	0.4768	100/0	0.00	
40	100	30.0	10.0	10.0	...	...	261	0.4767	100/0	0.00	10
40	100	30.0	10.0	10.0	...	...	197	0.4776	100/0	0.05	
40	100	30.0	10.0	10.0	$10^{-10}$	60	318	0.5291	77/23	0.00	11
50	100	30.0	10.0	10.0	$10^{-10}$	60	382	0.5573	93/7	0.00	12

is also consistent with the non-negligible mass accretion rate ( $10^{-10} M_{\odot} \text{ yr}^{-1}$ ).

However, whether modeled with a single WD component or a WD+disk composite model, the *FUSE* spectrum of EM Cyg reveals a rather hot WD. This result is in strong disagreement with the previous spectral analysis of (Welsh et al. 2005; Winter

& Sion 2003; Urban & Sion 2006) who found a WD temperature  $T \sim 20\text{--}24,000$  K. The discrepancy with previous analysis is due to (1) veiling of the WD by stream material overflowing the disk rim (in both the *FUSE* and *IUE* spectra around orbital phase  $\phi \sim 0.6\text{--}0.9$ , which were modeled as such); (2) contribution of the hot spot and secondary impact region to the longer

wavelengths of the *IUE* SWP spectra; and (3) eclipse of the disk in the SWP 08088 *IUE* spectrum (taken around orbital phase  $\phi \sim 0.98\text{--}0.16$ ). In addition, in these previous analysis, the flux was contributed mostly from an accretion disk, while in the present work the main component is the WD.

Originally, the derivation of EM Cyg system parameters were affected by the presence of the third star, *Arkadash*, in the background of the system. In a similar way, we find that the previously derived mass accretion rate and WD temperature might have been affected by the veiling of the WD and disk by stream overflow around orbital phase  $\phi \sim 0.65\text{--}0.90$ .

PG thank Steve (Stephen) Lubow for some interesting discussions on the stream–disk interaction, and Mario Livio for his kind hospitality at the Space Telescope Science Institute, where part of this work was carried out. We wish to thank an anonymous referee for her/his very prompt report and constructive criticism. We are thankful to the members of the American Association of Variable Star Observers (AAVSO) for providing public online optical archival data on EM Cyg. This research was based on observations made with the NASA-CNES-CSA Far Ultraviolet Spectroscopic Explorer. *FUSE* is operated for NASA by the Johns Hopkins University under NASA contract NAS5-32985. Support for this work was provided by the National Aeronautics and Space Administration (NASA) under Grant number NNX08AJ39G issued through the Office of Astrophysics Data Analysis Program (ADP) to Villanova University (P. Godon).

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